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Memorie della

High Performance Computing at INAF/OAPA

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Abstract. The scientific group at INAF/OAPA has a long experience in modeling astrophysical plasmas and in high performance computing (HPC). During 2000, INAF/OAPA has acquired a HPC facility, entirely devoted to extensive numerical modeling, in particular of astrophysical plasmas.

Key words. HPC – Hydrodynamics

1. Introduction: The Experience

1.1. Numerical Modeling

The effort for modeling coronal plasmas by the scientific group at INAF/OAPA has been stimulated by the X and UV solar missions since Skylab in 1973 (Vaiana et al. 1973; Vaiana & Tucker 1974).

Solar X-ray observations have shown that the emitting solar corona mostly consists of independent and isolated closed structures, the loops, where plasma at million degrees is confined by the coronal magnetic field. Fundamental questions to be answered then concern, first of all, the basic physics of the loop: what is the structure of the plasma inside them? why are they stable on time scales longer than their characteristic cooling times? what is the dynamics of the confined plasma? what is the heating source? Another interesting topic concerns, for instance, the investigation of highly transient and intense X-ray events, the coronal flares, occurring mostly inside closed loops.

In order to study coronal structures and phenomena, we start from the consideration that plasma confined in coronal loops mostly behaves as a fluid which moves and efficiently transports energy along the magnetic field lines. In such conditions, it is possible to describe the structure and evolution of the confined plasma by means of hydrodynamic models of compressible fluids, i.e. the solution of the equations of conservation of mass, momentum and energy, including the effects of thermal pressure, thermal conduction, which is very efficient in coronal conditions, compressional viscosity, solar gravity, radiative losses from optically thin plasma, local energy input which heats the plasma to coronal temperatures.

Even describing the system with only one coordinate, the one along the loop magnetic field lines, such models are often unfeasible for analytical treatment and require a numerical approach to be tackled. The group at INAF/OAPA has therefore developed a long experience in numerical hydrodynamic modeling. Hydrostatic loop

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modeling has been applied at first to derive scaling laws extended to loops higher than the pressure scale height and sustained by non-uniform heating (Serio et al. 1981), and then extensively to study the detailed structure of coronal loops and the diagnostics connected to solar and stellar X-ray observations (Giampapa et al. 1985; Peres et al. 1994; Maggio & Peres 1996; Di Matteo et al. 1998; Reale & Peres 2000a; Reale 2002; Testa et al. 2002). Steady hydrodynamic loop models have been applied to study the structure of siphon flows inside coronal loops (Orlando et al. 1995a,b) and the propagation of MHD Alfven waves in the solar wind (Orlando et al. 1996, 1997). Time-dependent hydrodynamic loop models, and in particular the Palermo-Harvard code (Peres et al. 1982; Betta et al. 1997), have been applied to study the stability of coronal loops (Peres et al. 1982), their heating by microflares (Peres et al. 1993; Reale et al. 1994; Betta et al. 1999), their ignition (Reale et al. 2000b,c), the evolution of solar coronal flares (Peres et al. 1987; Serio et al. 1991; Jakimiec et al. 1992; Peres & Reale 1993a,b; Reale & Peres 1995; Reale et al. 1997; Betta et al. 2001), the diagnostics of stellar X-ray flares (Reale et al. 1988, 1993; Reale & Micela 1998).

The interest has also extended to study astrophysical systems with non-confined plasmas which cannot be described with 1-D models. The thermal stability of plasmas in stratified atmospheres (Reale et al. 1991, 1994, 1996) has been studied by means of 2-D time-dependent hydrodynamic models, based on an improved Flux-Corrected-Transport (FCT) technique (Reale et al. 1990a), and on a time-splitted Alternating-Direction Implicit (ADI) numerical scheme to describe thermal conduction (Reale 1995). Preliminary studies have been conducted on the interaction of SNR with the interstellar medium (Maggio et al. 1994; Orlando et al. 2002). The same 2-D model has been applied to study the propagation of Coronal Mass Ejections (Ciaravella et al. 2001) and the evolution of flares triggered in a non-confined corona (Reale et al. 2002).

1.2. HPC experience

Many of the problems described above have required extensive numerical resources and high performace computing. A vector version of the Palermo-Harvard code has been executed on the Cray systems at CINECA (Bologna) (Reale et al. 1988). The group has also acquired pioneeristic experience on the development and optimization of parallel codes on Transputer networks for solving astrophysical problems, both based on PC's in collaboration with IAIF/CNR (Palermo) (Reale et al. 1990b), and on self-standing Unix-based systems, and in particular a Meiko Computing Surface at INAF/OAPA (Reale 1990; Reale et al. 1992). Parallel codes have then been ported and tested for parallel efficiency on Unix-based general-purpose Local Area Networks (LAN), using Fortran codes and publicly available parallel libraries, such as Parallel Virtual Machine (PVM) (Reale et al. 1994b). A high level library for parallelization of Fortran programs tailored for domain-decomposition approach has been developed in-house (Reale et al. 1992). The 2-D FCT/ADI code has been also ported for efficient execution on the Cray T3E by means of the HPF parallel library (Reale & Bocchino 1999).

More recently, INAF/OAPA has had the opportunity to build up a HPC facility, named "Sistema di Calcolo per l'Astrofisica Numerica" (SCAN, Computing System for Numerical Astrophysics), and entirely devoted to massive numerical computing in Astrophysics.

2. SCAN

2.1. The project and the hardware selection

SCAN has been built in the framework of a proposal named "High education in

HPC and current Astrophysical problems" approved by Ministero dell'Universitá e della Ricerca Scientifica e Tecnologica and funded by European Union in 1999.

The project involved essentially two phases:

- High education of young personnel, hired through specific selection procedures, and trained through study, advanced research and technical managing activity.
- Hardware set up, including the settlement of a HPC system, selected through a competitive inquirement of specialized firms, its installation and testing, and its production activity on specific numerical problems.

For the hardware selection, we have inquired four firms, specialized for HPC systems on the required scale-size: IBM, SUN, Compaq and SGI. The firms were asked to propose their best HPC system within a set of constraints and a given budget scale, and to perform a benchmark program on that system.

The constraints were to propose a self-standing and complete system, Unix-based, equipped with standard parallel software (e.g. HPF, MPI), maximized for computing speed and with a RAM $\geq$ 512 MB per processor.

The benchmark consisted in the execution of the 2-D FCT hydrodynamic code described in section 1.1 on a specific astrophysical problem. The problem is the free propagation of a Coronal Mass Ejection (CME) expelled from the solar surface (Ciaravella et al. 2001). It is described the evolution for 500 s of an isothermal corona (1.5 MK) with an isobaric, dense and circular plasma blob, ejected upwards at 400 km/s from the low layers of the atmosphere and thermally insulated from the surrounding medium. The 2-D hydrodynamic equations of conservation of mass, momentum and energy are solved on an 800×800 grid, without thermal conduction. The code version is written in F77 and parallelized with HPF (Reale & Bocchino 1999) for a total of $\sim 2.4 \times 10^{11}$ floating point operations. This benchmark problem is a typical medium-size astrophysical plasma problem. The code parallelization is based on a plain domain-decomposition approach: each processor makes the identical computations on a different slice (with the same size) of the geometrical domain. The FCT technique requires a limited exchange of values at the boundaries of adjacent slices (message-passing). For the given grid size, we expect a good scaling of the parallel performances for a medium-grain parallel system, with about ten processors.

Fig.1 shows the evolution of the density and of the velocity field throughout the 500 s of computations according to the benchmark problem. The density contrast over the unperturbed medium is shown, in order not to display the plasma gravitational stratification. It is clearly visible the upward propagation of the blob, initially close to the lower layers of the atmosphere. While it propagates, it expands adiabatically and a bow shock develops and propagates ahead of the blob. Over the 500 s time lapse the overall blob velocity has not have enough time to be reduced significantly from the initial value by the gravitational braking.

All firms have proposed one or more alternative systems and have accepted to execute the benchmark program, providing results for different configurations and numbers of CPUs. Fig.2 shows the performances obtained with various systems proposed by the firms, in terms of speed in MFlops vs number of processors simultaneously used.

After a careful evaluation of the benchmark results and of the other aspects of the proposals of each firm, the administration of OAPA ended up to select the computing system proposed by the Compaq firm for equipping SCAN. The system has been acquired and installed in Summer 2000, and after an initial period of tuning up and testing, it has been running for production continuously.

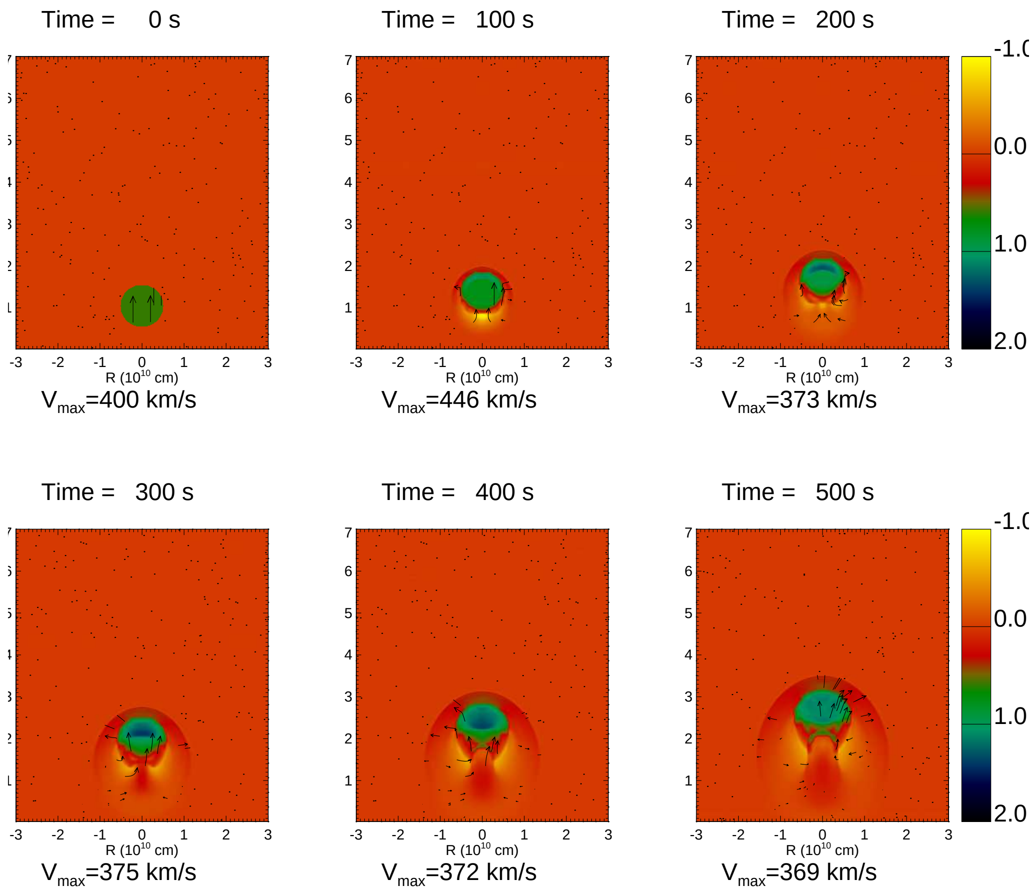

Fig. 1. Evolution of the density contrast (over the unperturbed atmosphere) of a blob of coronal plasma ejected upwards at 400 km/s from the solar surface computed with the 2-D FCT hydrodynamic code. This is the benchmark problem set up for the selection of a computing system tailored for the HPC program at INAF/OAPA. The arrows mark the velocity field, normalized to the maximum speed labelled under each frame.

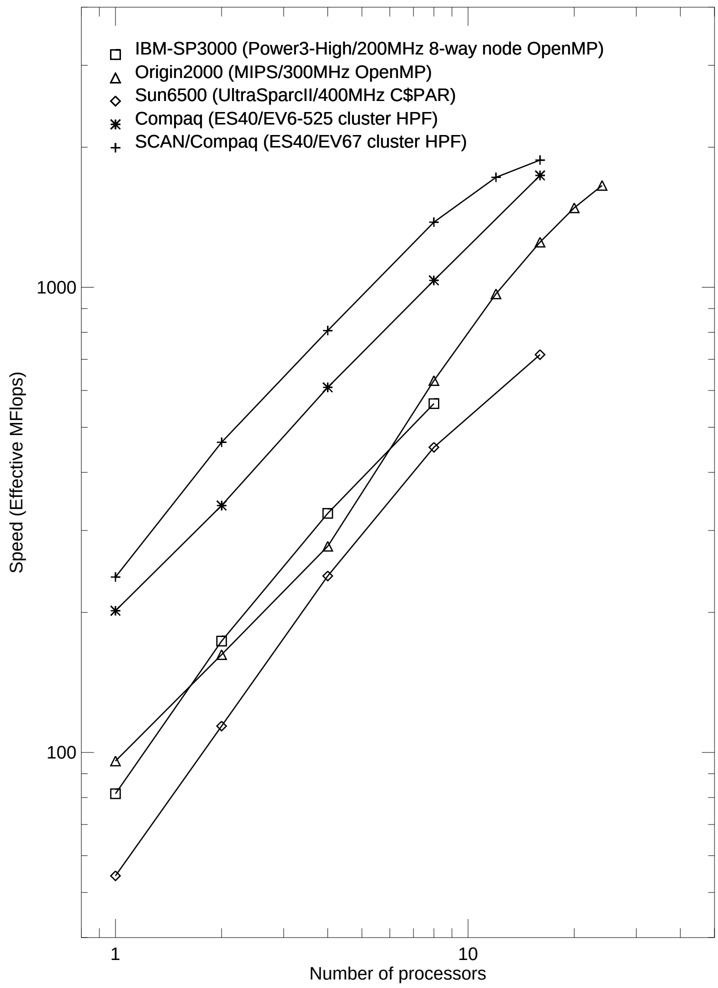

Fig. 2. Performances of the benchmark problem obtained with the labelled computing systems for different numbers of processors. Each curve and symbol pertain to a computing system. Best performances (upper curve) are obtained with the actually selected Compaq Cluster equipped with EV67 processors.

2.2. Description

The HPC system at SCAN consists of a Unix-based COMPAQ HPC cluster, including four AlphaServers ES40, each with four EV67 (667 MHz) processors, for a total of 16 processors, interconnected with an efficient Memory Channel II device. With the newest versions of the operating system, it is a single-image system, i.e. all servers are managed together and seen as a single computer by the operating system.

The present configuration of SCAN is shown in Fig.3. The HPC system is placed in a thermally and acustically insulated chamber inside a larger room which includes several working places for computing personnel. The room is located in a separate section of INAF/OAPA including also the XACT (X-ray Astronomy Calibration and Testing) laboratory, distant $\sim$ 500 m from the main INAF/OAPA building.

The SCAN HPC system has its own LAN with several Linux PCs for direct connection to the mainframe. The SCAN LAN is connected to the INAF/OAPA LAN through a dedicated router and an optical bridge realizes the physical link between the SCAN/XACT building and the main INAF/OAPA building.

2.3. Scientific applications

The SCAN facility has already produced scientific results worth of publication on specialized journals. An extensive set of hydrodynamic simulations have been executed on the HPC system of SCAN in the framework of the study of the dynamics and diagnostics of flares occurring in non-confined coronae (Reale et al. 2002). The motivation for this modeling effort stems from the observation of stellar flares which are often much longer and more intense than typical solar flares. In such conditions it is legitimate to wonder whether the assumption of plasma confinement typically valid for solar flares still holds for such large scale stellar flares or not. A way to explore this question is to model hypothetical flares occurring in a non-confined atmospheres and to derive characteristic signatures of non-confinement vs confinement. Non-confined flares have been modeled as triggered by a heating pulse as already done for confined flares. Is used the time-dependent 2-D FCT/ADI hydrodynamic code described in Sec. 1.1. Computations are done on an inhomogeneous but fixed numerical grid in cylindrical R-Z coordinates. Preliminary simulations have been performed on the CRAY T3E at CINECA, and the bulk of an extensive set of production simulations on the SCAN HPC system. This set of simulations program aimed at the exploration of the space of the relevant physical parameters, i.e. the position, intensity and duration of the heating pulse, and the pressure of the unperturbed atmosphere.

Fig. 4 shows an example of the computed evolution of the non-confined plasma subject to an intense heating pulse. At time t=0 the heating is abrutply switched on and a thermal conduction front immediately develops propagates isotropically from the heating site (at the center of the bright region in the upper right frame). This front hits against the dense chromosphere at the base of the corona, which is suddenly heated and expands upwards, determining the formation of a shell-like evaporation front, well visible in the middle left panel. This hot density front determines the X-ray flare. As soon as the heating stops, at t = 50 s, the temperature suddenly drops but the density front still propagates for some time, as shown in the lower panels.

One result of this study is that, after the heating stops, the decay of the X-ray emission invariably occurs on very short time scales, few minutes, much shorter than those of observed stellar flares. This is mostly due to the very fast cooling of the plasma. This is a strong indication that large-size stellar flares still occur in confined plasma configurations.

More recently INAF/OAPA has started a collaboration with the ASCI/FLASH center, for extension and application of the advanced FLASH code to coronal plasma problems, as described in detail elsewhere in this issue (Orlando et al. 2002). In this context the group at INAF/OAPA has contributed to extend the FLASH code with modules for computing plasma thermal conduction, radiation from optically thin plasma and the fractionation of ion species in non-equilibrium of ionization. Preliminary applications of the FLASH code on coronal plasma problems at INAF/SCAN include the simulation of the initial phases of a soft X-ray flare and the interaction of a SNR shock front with a dense perturbation of the interstellar medium both including the effects of non-equilibrium of ionization.

3. Conclusions

OAPA/SCAN is a facility entirely devoted to massive numerical computing. It is tailored and scaled for computing ef-

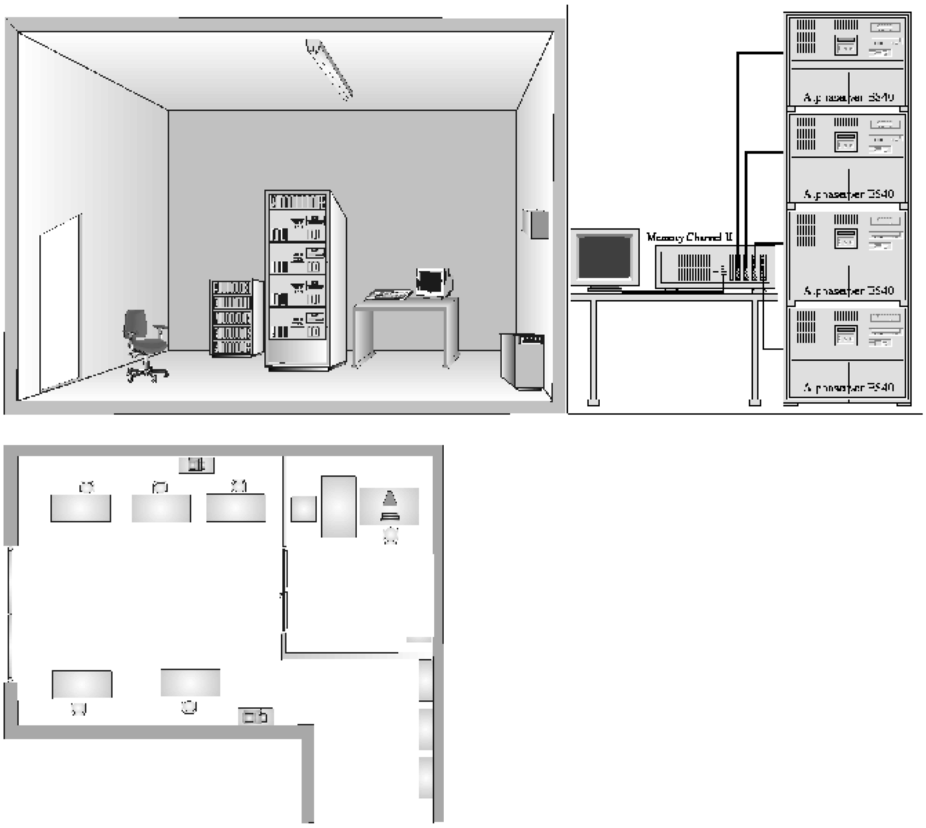

Fig. 3. Actual configuration of the SCAN facility. The HPC system (upper right panel) is disposed in a thermally and acoustically insulated chamber (upper left panel) in turn located in a larger room containing several working places (lower panel).

ficiently plasma 2-D hydrodynamics, but it is a geneal-purpose system and can be safely applied to many other computational modeling projects. Numerical resources of SCAN are available to the INAF community.

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References

Betta, R. M., Peres, G., Serio, S., Reale, F., 1997, A&AS, 122, 585.

Betta, R., Reale, F., Peres, G., 1999, in Plasma Dynamics and Diagnostics

Betta, R., Peres, G., Reale, F., Serio, S., 2001, A&A, 380, 341. in the Solar Transition Region and Corona, ESA SP 446, 179

Ciaravella, A., Raymond, J.C., Reale, F., Strachan, L., Peres, G., 2001, ApJ, 557, 351.

Di Matteo, V., Reale, F., Peres, G., Golub, L. , 1998, A&A, 342, 563.

Giampapa, M.S., Golub, L., Peres, G., Serio, S., Vaiana, G.S., 1985, ApJ, 289, 203.

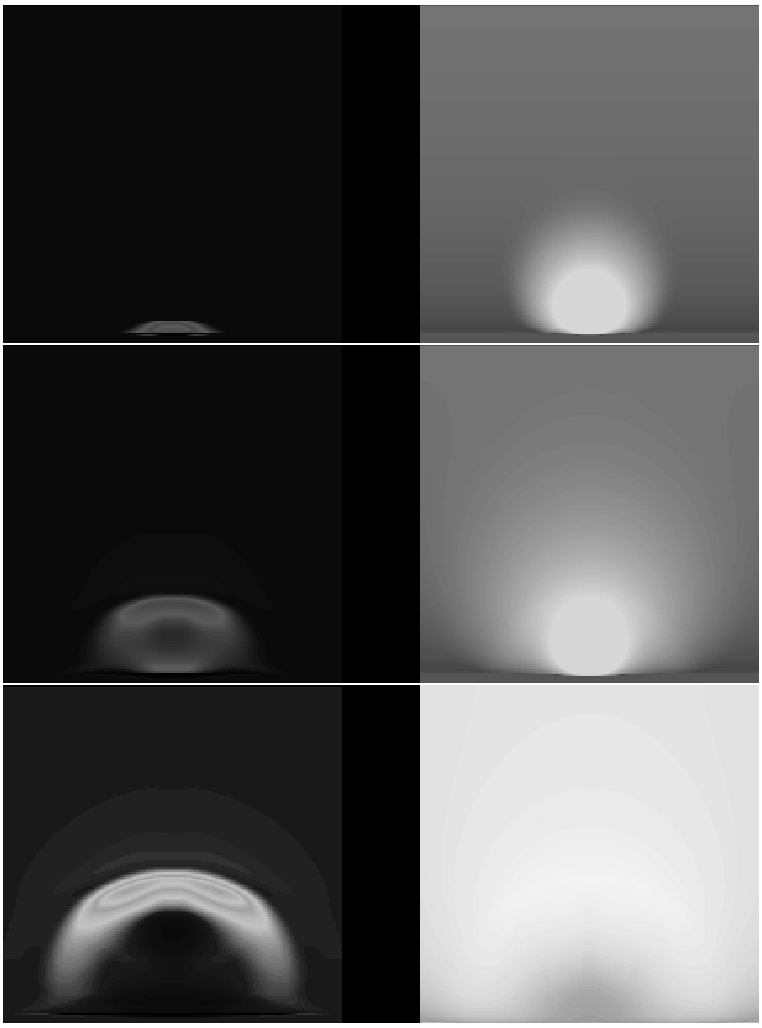

Fig. 4. Three snapshots during the evolution of a flare occurring in a non-confined corona: slices of density (*left*), temperature (*right*) at time 10, 50 and 100 s since the beginning of the simulation.

Jakimiec, J., Sylwester, B., Sylwester, J., et al. 1992, A&A, 253, 269

Maggio, A., Reale, F., Peres, G., Ciaravella, A., 1994, Comp. Phys. Comm., 81, 105

Maggio, A., Peres G., 1996, A&A, 306, 553.

Orlando, S., Peres, G., Serio, S., 1995a, A&A, 294, 861.

Orlando, S., Peres, G., Serio, S., 1995a, A&A, 300, 549.

Orlando, S., Lou, Y.-Q., Rosner, R., Peres, G., 1996, JGR, 101, 24443.

Orlando, S., Lou, Y.-Q., Rosner, R., Peres, G., 1996, JGR, 102, 24139.

Orlando, S., Peres, G., Reale, F., Rosner, R., Plewa, T., Siegel, A., this issue.

Peres, G., Rosner, R., Serio, S., Vaiana, G. S., 1982, ApJ, 252, 791.

Peres, G., Reale, F., Serio, S., Pallavicini, R., 1987, ApJ, 312, 895.

Peres G., Reale F., Serio S., 1993, in Physics of Solar and Coronae, J.L. Linsky, S. Serio (eds.), Kluwer Academic Pub., 151.

Peres, G., Reale, F., 1993, A&A, 566, 267.

Peres, G., Reale, F., 1993, A&AL, 275, L13.

Peres G., Reale F., Golub L. , 1994, ApJ 422, 412
Reale, F., Peres, G., Serio, S., Rosner, R., Schmitt, J.H.M.M., 1988, ApJ, 328, 256.
Reale, F., Peres, G., Serio, S., 1990, Nuovo Cimento B 105, 1235.
Reale, F., Brugè, F., Peres, G., Fornili, S.L., Martorana, V., Serio, S., 1990b, Comp. Phys. Comm., 201, 60.
Reale, F., 1990, Parallel Computing, 16, 361.
Reale, F., Rosner, R., Malagoli, A., Peres, G., Serio, S., 1991, MNRAS, 251, 379.
Reale, F., Barbera, M., Sciortino, S., 1992, Comp. Phys. Comm. 72, 129
Reale, F., Serio, S., Peres, G., 1993, A&A, 272, 486.
Reale, F., Bocchino, F., Sciortino, S., 1994, Comp. Phys. Comm. 83, 130.
Reale, F., Serio, S., Peres, G., 1994, ApJ, 433, 811
Reale, F., 1995, Comp. Phys. Comm., 86, 13.
Reale, F., Peres, G., 1995, A&A, 299, 225.
Reale, F., Peres, G., Serio, S., 1996, A&A, 316, 215
Reale, F., Micela, G., 1998, A&A, 334, 1028.
Reale, F., Betta, R., Peres, G., Serio, S., McTiernan, J., 1997, A&A, 325, 782
Reale, F., Bocchino, F., 1999, Proc. Fifth European SGI/Cray MPP Workshop, BOLOGNA (Italy) - September 9-10, 1999, http://www.cineca.it/mpp-workshop/proceedings.htm
Reale, F., Peres, G., 2000, ApJL, 527, L45.
Reale, F., 2002, ApJ, 580, 566.
Reale, F., Bocchino, F., Peres, G., 2002, A&A, 383, 952.
Reale, F., Peres, G., Serio, S., DeLuca, E.E., Golub, L., 2000, ApJ, 535, 412.
Reale, F., Peres, G., Serio, S., Betta, R., DeLuca, E.E., Golub, L., 2000, ApJ, 535, 423.
Serio, S., Peres, G., Vaiana, G.S., Golub, L., Rosner, R., 1981, ApJ 243, 288.
Serio, S., Reale, F., Jakimiec, J., Sylwester, B., Sylwester, J., 1991, A&A, 241, 197.
Testa, P., Peres, G., Reale, F., Orlando, S., 2002, ApJ, 580, 1159.
Vaiana, G.S., Davis, J. M., Giacconi, R., Krieger, A.S., Silk, J.K., Timothy, A.F., Zombeck, M., 1973, ApJL, 185, L47
Vaiana, G.S., Tucker, W., 1974, in X-ray Astronomy, eds. R. Giacconi, and H. Gursky, Dordrecht:Reidel